

# NIR spectroscopy of the most massive open cluster in the Galaxy: Westerlund 1

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**Abstract.** Using ISAAC/VLT, we have obtained individual spectra of all NIR-bright stars in the central 2'x2' of the cluster Westerlund 1 (Wd 1) with a resolution of  $R \approx 9000$  at a central wavelength of  $2.30 \mu\text{m}$ . This allowed us to determine radial velocities of ten post-main-sequence stars, and from these values a velocity dispersion. Assuming virial equilibrium, the dispersion of  $\sigma = 8.4 \text{ km/s}$  leads to a total dynamical cluster mass of  $1.25 \times 10^5 M_\odot$ , comparable to the photometric mass of the cluster. There is no extra-virial motion which would have to be interpreted as a signature of cluster expansion or dissolution.

**Keywords.** open clusters and associations: individual: Westerlund 1 – Galaxies: star clusters – supergiants

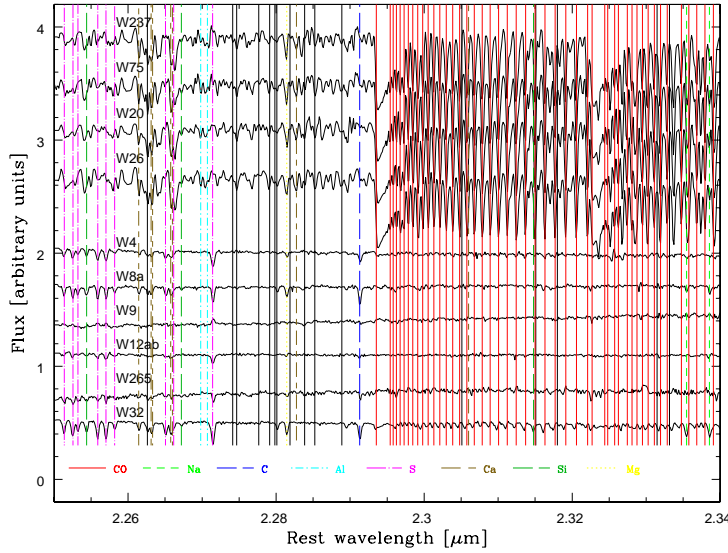
## 1. Observations and Results

We used ISAAC/VLT in its highest spectral resolution mode ( $R \approx 9000$ ) to step the 2' long slit across the centre of the cluster for 2', in steps of  $0.3''$ . At each position, the integration time was  $3 \times 1.77 \text{ s}$ . The central wavelength of  $2.31 \mu\text{m}$  was optimized for the analysis of red supergiant (RSG) spectra, which show deep CO absorption features in this regime. The data were sky subtracted, flat fielded, distortion corrected and wavelength calibrated using standard *iraf* routines and, where necessary, calibration observations. Individual spectra were extracted from those slit positions where a star with photospheric features in our wavelength range was included. After the extraction of the 1d spectrum, we corrected for telluric absorption. Fig. 1 shows the ten spectra which were extracted from our scan, four red supergiants (RSGs), five yellow hypergiants (YHG), and one sgB[e] star.

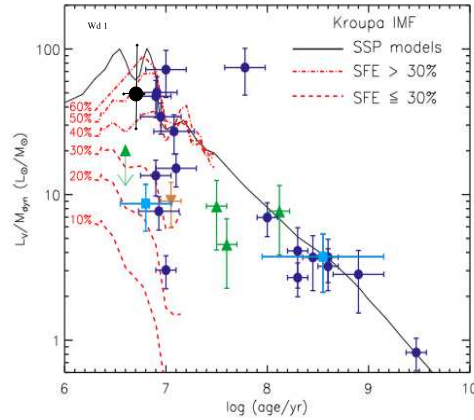
We used the lines in these spectra to determine the radial velocities and their dispersion, which is  $\sigma = 8.4 \text{ km/s}$ . This, together with the half-light radius (estimated from archival NTT/SOFI images), assuming that the Virial Theorem is valid, yields a dynamical cluster mass of  $1.3 \times 10^5 M_\odot$  (see also Mengel & Tacconi-Garman (2007)).

Comparing the light-to-mass-ratio (V-band luminosity and dynamical mass) to that expected for a cluster of this age and luminosity, and a Kroupa IMF, solar metallicity, allows a conclusion regarding the dynamical state of the cluster: Even though it could in principle also be an indication of an excess of low-mass stars (compared to the Kroupa IMF), the more likely explanation for low L/M is that the high “dynamical mass” arises from extra-virial motion, and hence an expanding cluster. Fig. 2 shows a plot taken from Goodwin & Bastian (2006), where they determine the velocity dispersion of the stars in their N-body simulated clusters and plot the resulting L/M ratios as dashed lines. Clusters with a low star-formation efficiency show large amounts of extra-virial motion, and do not survive as bound clusters.

The L/M ratio for Westerlund 1 is consistent with the expectation from the evolution-



**Figure 1.** Normalized, vertically offset spectra of all detected cluster members with absorption features in our observed wavelength regime. Top four are the RSGs, five of the six below are YHG (W9 is a sgB[e] star).



**Figure 2.**  $L_V/M_{dyn}$  as a function of age, for “static” clusters (solid) and clusters adjusting dynamically after gas expulsion (from Goodwin & Bastian (2006)), dashed. Wd 1 (black dot) does not seem to be expanding dramatically.

ary synthesis model, which means that the cluster may survive for a long time, possibly several Gyr.

The most important caveat in our measurement is the low number of stars, which is why we want to improve statistics by obtaining radial velocities also for the hot stars in the cluster, from absorption lines around Br $\gamma$ .

## References

- Goodwin S.P & Bastian N. 2006, *MNRAS* 373, 752  
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